

Wide-field aurora imager onboard Fengyun satellite: Data products and validation

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Key Points:

- The WAI can image the aurora with wide FOV, high sensitivity, and hightemporal and spatial resolutions in FUV wavelengths.
- The aurora data from WAI agree well with the SSUSI data in both aurora oval distribution and boundaries.
- The WAI observations can describe the aurora oval evolution with upstream solar wind parameters in detail.

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Abstract: New observations of auroras based on the wide-field aurora imager (WAI) onboard Fengyun-3D (FY-3D) satellite are exhibited in this paper. Validity of the WAI data is analyzed by comparing auroral boundaries derived from WAI observations with results obtained from data collected by the Special Sensor Ultraviolet Spectrographic Imager (SSUSI) aboard the Defense Meteorological Satellite Program (DMSP F18). Dynamic variations of the aurora with the solar wind, interplanetary magnetic field (IMF) parameters, and the SYM-H index are also investigated. The comparison of auroral boundaries indicates that the WAI data are morphologically valid and suitable to the study of auroral dynamics. Effective responses to solar wind parameters indicate that the WAI data can be useful to monitor and predict the Earth's space weather. Since the configuration of aurora is a good indicator of the solar wind–magnetosphere–ionosphere (SW-M-I) coupling system, and can reflect the disturbance of the space environment, the WAI will provide important data to help us to study the physical processes in space.

Keywords: WAI; FY-3D; auroral dynamics; SSUSI; SW-M-I

1. Introduction

Auroras are produced by collisions between energetic magnetospheric particles precipitating into the polar region through the Earth's magnetic field lines and the atoms and molecules in the upper atmosphere (Meier, 1991). Variations of aurora are closely related to the dynamics of the ionosphere, the magnetosphere, and the solar wind, and the couplings between them. The configuration of the aurora is a good indicator of the solar wind–magnetosphere–ionosphere (SW-M-I) coupling system. Also, the aurora is highly dynamic during storms and substorms (Germany et al., 1998). Studying aurora is a valid method to explore the dynamics processes in the magnetosphere and the ionosphere.

Studying the aurora phenomenon can help us to profoundly understand the processes in the solar–terrestrial space. For example, Liou and Sotirelis used auroral data from the Ultraviolet Imager (UVI) on the POLAR satellite to study responses of the northern winter polar cap to 28 auroral substorms; they found that the auroral breakup during substorms is preceded by inflation of the polar cap and followed by deflation of the polar cap (Liou and Sotirelis, 2016). Milan utilized the auroral images obtained by the Wide-band Imaging Camera (WIC) onboard the IMAGE satellite to track changes in the polar cap area and to quantify the amount of open flux in the magnetosphere (Milan et al., 2012). Zhang and Paxton used auroral data from the Global Ultraviolet Imager (GUVI) onboard the TIMED satellite to construct an empirical K_p -dependent global auroral model (Zhang and Paxton, 2008). Luan utilized SSUSI/DMSP data to study the dependences of the auroral energy flux distribution over the nightside oval with seasonal variety, geomagnetic parameters, and solar activity (Luan XL et al., 2018). Ding used GUVI/TIMED and SSUSI/DMSP auroral data to develop

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an automatic auroral boundary determination algorithm (Ding GX et al., 2017).

The observation of auroras has gone through a long period. Auroral configuration has been described by unaided human visual observation, and by instrumental observation from ground-based stations, and more recently via imaging from space. Imaging the auroral oval from space provides spatial/temporal information of the precipitation particles that allows these particles to be mapped back to different source regions in the magnetosphere, which allows us to understand the energy transportation processes (Germany et al., 1994a, b). Since the 20th century, more and more satellites have been sent to space to obtain more precise spatial data; the aurora is especially important among the observed phenomena. For example, the UVI/POLAR, the WIC/IMAGE, the GUVI/TIMED, and the SSUSI/DMSP (Paxton et al., 2004; Sotirelis et al., 2013). The WAI, developed by the Institute of Optics, Fine Mechanics, and Physics, Chinese Academy of Sciences, is aboard the FY-3D satellite, which was launched on 15 November 2017; on 8 December 2017, it began routinely acquiring scientific data. Since then, the ground application system of FY-3D has accumulated a large number of auroral images.

Both the GUVI and SSUSI spectrographic instruments are of the scanning type, with limited Field of View (FOV); they cannot capture the global auroral ovals due to the low altitude of their sun-synchronous satellite orbits (Paxton et al., 1992, 1993, 2004). The WAI has a novel design that allows it to get relatively wide-area auroral images at low altitudes, compared to those from the SSUSI. The purpose of this paper is to verify the validity of WAI data by comparing auroral configurations derived from WAI data with the observations of SSUSI, and by investigating the auroral evolution with variable solar wind. The system design, instrument development, and the data processing algorithm that transfer photon signals to disk images projected in geographic/geomagnetic grids on the reference sphere of WAI have been introduced in Zhang's paper (Zhang XX et al., 2019).

2. Materials and Methods

2.1 Data Sets

The auroral images used in this study were obtained by the WAI/FY-3D and the SSUSI/DMSP F18. The FY-3D satellite operates in a sun-synchronous orbit with an orbital height of about 836 km. The instantaneous FOV of the WAI optical system is $130^\circ \times 10^\circ$ (130° corresponds to the cross-track direction and 10° corresponds to the along-track direction) with operational wavelength range 140–180 nm. In scanning mode, the WAI can achieve a total FOV of $130^\circ \times 130^\circ$, which allows relatively complete auroral images to be obtained. Since the launch of the FY-3D on 15 November 2017, the WAI has accumulated and returned a large number of auroral images. The WAI data used in this paper are provided by the ground application system of the FY-3D. The data processing algorithm operated in the ground application system is described in Zhang's paper (Zhang XX et al., 2019).

Another category of auroral images, used here as a comparison for the images obtained from WAI data, were obtained by the SSUSI instruments (Paxton et al., 2004; Sotirelis et al., 2013). The

DMSP satellites were launched into circular polar orbits with orbital height of 840 km. From F16 through F18, the DMSP satellites were equipped with SSUSI; SSUSI data are available from 2004 on. The SSUSI optical system includes an imaging spectrograph with instantaneous FOV of 11.8° , with five channels whose wavelengths are between 115 and 180 nm (including the LBH band from 140 to 180 nm), and equipped with cross track scanning mode. The spectrographic imager was designed to provide quantitative observations and interpretations of the Earth's airglow and auroral emissions in FUV. The SSUSI auroral images used in this paper were obtained from the websites of <http://ssusi.jhuapl.edu/>.

The solar wind and interplanetary magnetic field (IMF) parameters used in this paper were provided by NASA OMNI Web. The three components of IMF, the solar wind density and velocity, and the solar wind dynamic pressure, are in 1 min resolution. Since the data were obtained at the nose of the magnetopause, it is necessary to consider the delay that is known to occur as a phenomenon propagates to the ionosphere. The propagation times are estimated to be the sum of 5 min for traversing the magnetosheath and 2 min Alfvén transit time to reach the ionosphere (Liou et al., 1998).

2.2 Comparison with DMSP

The SSUSI auroral data have been proverbial in studies of space dynamics transversion. In this section, we perform detailed comparisons between SSUSI observations and WAI observations in order to establish the validity and accuracy of the WAI results. These comparisons include auroral shapes and boundaries.

To examine the consistency of the auroral radiation distribution between WAI and SSUSI, we compared the adjoining time observations of the two instruments. The geomagnetic activity was basically quiet on 30 August 2018, with a maximum K_p of 2. The auroral observations at about 00:15 (UT) are shown in Figure 1. The red asterisks in (a) and black asterisks in (b) show the equatorward boundaries that were extracted with the fuzzy local information C-means clustering segmentation algorithm (FLICM) (Yang QJ et al., 2013; Ding GX et al., 2017). The FLICM algorithm incorporates local spatial and gray level information in a fuzzy way in order to preserve robustness and noise insensitiveness, by adding a novel fuzzy factor

$$G_{ki} = \sum_{\substack{j \in N_i \\ i \neq j}} \frac{1}{d_{ij} + 1} (1 - u_{kj})^m \|x_j - v_k\|^2 \quad (1)$$

to the squared error objective function

$$J_m = \sum_{i=1}^N \sum_{k=1}^c [u_{ki}^m \|x_i - v_k\|^2 + G_{ki}], \quad (2)$$

where N is the total number of pixels; c is the number of clusters with $2 \leq c < N$ (in this paper, we set c to be 6); x_i is the gray level of the i^{th} pixel that is the center of the local window (e.g., 3×3 pixels); v_k is the prototype of the center of cluster k ; u_{ki} is the degree of the membership of x_i in the k^{th} cluster; m is a weighting exponent on each fuzzy membership that determines the amount of fuzziness of the resulting classification, and is usually set to be 2

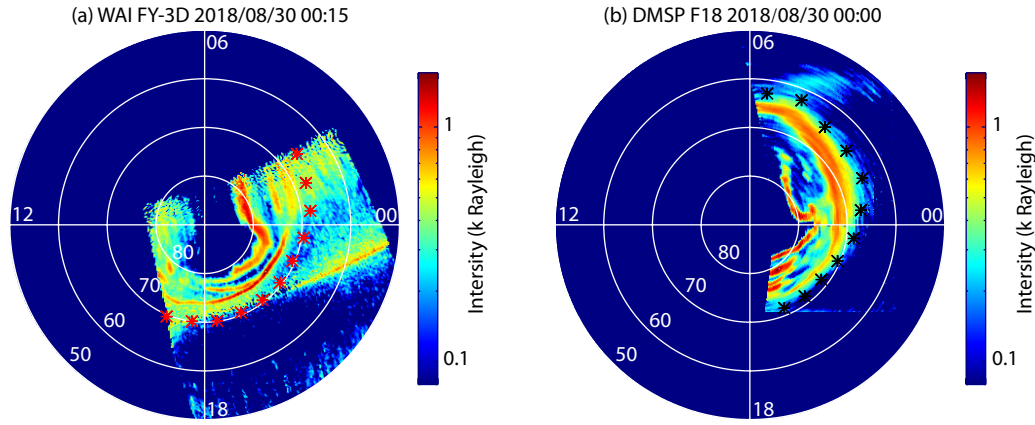


Figure 1. Aurora images observed by (a) FY-3D WAI and (b) DMSP F18 SSUSI in the southern hemisphere at 00:15 and 00:00 on 30 August 2018. Two images have a time interval of 15 minutes. The red asterisks in (a) and black asterisks in (b) are the equatorward boundaries extracted by the FLICM algorithm. Midnight is on the right and noon is on the left.

(Bezdek, 1981); N_i stands for the set of neighbors falling into a local window around pixel x_i ; and d_{ij} is the spatial Euclidean distance between pixels i and j . The accuracy of the auroral boundaries extracted by FLICM has been verified in several studies (Yang QJ et al., 2013; Ding GX et al., 2017). It can be seen from Figure 1 that the auroral boundaries of WAI origin are quite similar to those from SSUSI data.

The comparison of auroral images from WAI and SSUSI (given in Figure 2) testifies to the consistency of their detection of auroral morphology. In order to eliminate differences between the outputs of two instruments, normalization of data between the two instruments was done before merging. For the locations that were observed by both WAI and SSUSI, the average intensities were calculated for each, I_{WAI} for WAI and I_{SSUSI} for SSUSI, respectively. Then the normalization factor I_{SSUSI}/I_{WAI} was used to normalize the

WAI data to the level of SSUSI; that is, all the pixels in the WAI image are multiplied by I_{SSUSI}/I_{WAI} . The red asterisks in Figure 2 are the auroral equatorward boundaries derived from WAI observations; the black asterisks are the auroral equatorward boundaries derived from SSUSI data. Figure 2 demonstrates that the auroral morphologies of WAI and SSUSI are in good agreement. The auroral radiation distribution tendencies and the local features of the two observations also show quite fine agreement. The locations of the equatorward boundaries are basically identical. These indicate that the WAI observations are highly consistent with those of the SSUSI.

A median geomagnetic storm, with the maximum Kp of 6, erupted between 5–7 May, 2018. The observations of auroral ovals at about 20:44 UT, 5 May 2018 are shown in Figure 3, in which the red and black asterisks are the equatorial boundaries extracted with the FLICM algorithm. Figure 4 superimposes the auroral images of Figure 3. It can also be seen from the figures that the distribution tendencies of the auroral radiation and the local features of the two observations show quite fine agreement. From the two cases, the trends of auroral boundaries and the similarity of the featured regions can demonstrate that the observations of WAI have fine consistency with those of SSUSI.

Furthermore, in order to verify the match between auroral images from WAI and observations from SSUSI, southern hemisphere auroral images from June 15, 2018 to June 24, 2018 that were observed by both WAI and SSUSI were selected for statistical analysis. Auroral images from WAI during this time period were matched one at a time with the corresponding images from SSUSI. A time window of 10 min was chosen to select the auroral images from the two instruments to minimize temporal evolution of the auroral oval between the two observations. We have assumed that the auroral oval is sufficiently stable for meaningful comparison at the 10 min time scale; however, this may not be true of observations made during the main phase of storms or the expansion phase of substorms. Once an image pair satisfying this criterion was determined, the two data sets were mapped into the same grid of magnetic latitude (MLAT) and magnetic local time (MLT) in the corrected (altitude-adjusted) geomagnetic coordi-

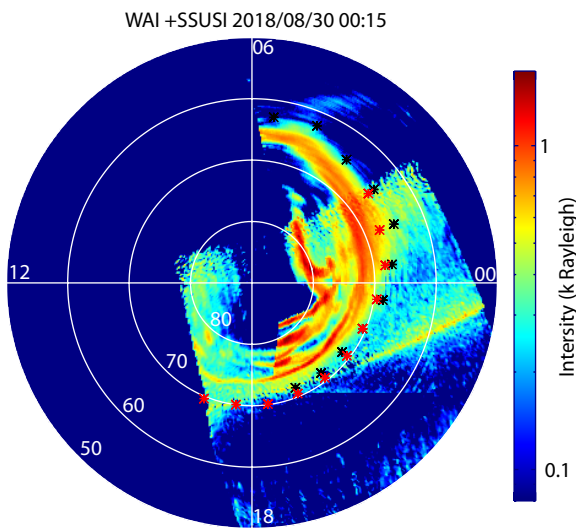


Figure 2. The superposition of aurora images derived from WAI and from SSUSI observations. The superposed auroral images come from Figure 1. The normalization factor I_{SSUSI}/I_{WAI} has been used to normalize the WAI data to the level of SSUSI data. The red and black asterisks are the equatorward boundaries.

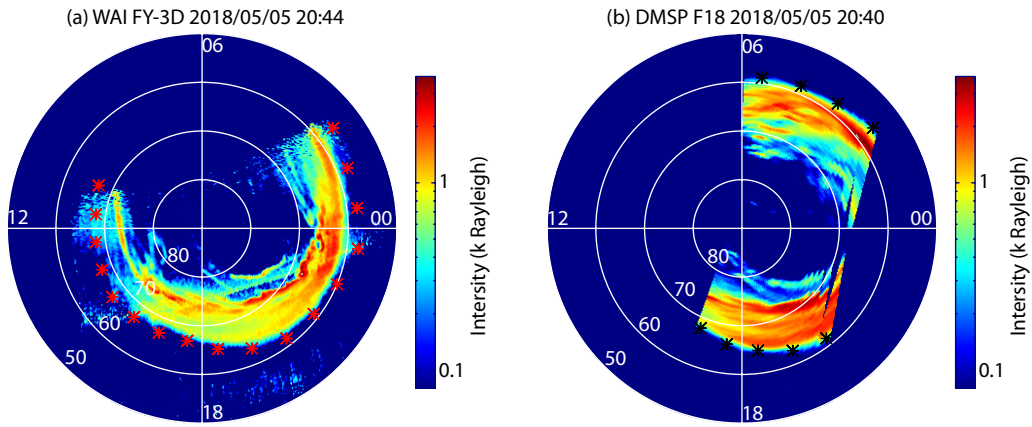


Figure 3. Aurora images observed by FY-3D WAI and DMSP F18 SSUSI in the southern hemisphere at 20:44 and 20:40 on 5 May 2018. The annotations are same as for Figure 1.

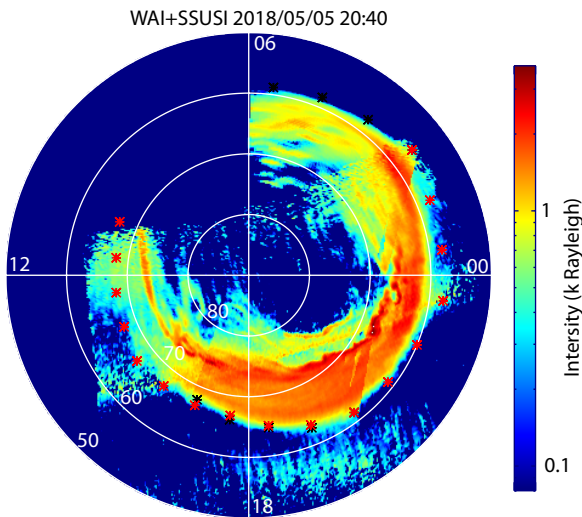


Figure 4. Superimposition of WAI and SSUSI aurora images from Figure 3. The annotations are the same as for Figure 2.

ates. Then the auroral equatorward boundaries were extracted from the two data sets, using the FLICM algorithm. The statistics of two kinds of boundary points are shown in Figure 5. As can be seen, there are 347 pairs of boundary points in total and the correlation coefficient is 0.927. All the scatters are linearly fitted to straight line of $MLAT_{SSUSI} = A + B \times MLAT_{WAI}$ as shown in the lower part in Figure 5. A is 1.310, B is 0.983 and the uncertainties of A and B are 1.300 and 0.020. The gray solid line indicates $MLAT_{SSUSI} = MLAT_{WAI}$, that is, $MLAT_{SSUSI}$ and $MLAT_{WAI}$ agree perfectly; the dashed lines are plotted at $MLAT_{SSUSI} = MLAT_{WAI} \pm 3^\circ$. The standard deviations between $MLAT_{SSUSI}$ and $MLAT_{WAI}$ is 1.038°, and 68.876%, 98.559%, 98.847% of the scatters within $\pm 1^\circ$, $\pm 2^\circ$, $\pm 3^\circ$, respectively. In summary, the boundaries derived from WAI observations agree very well with those from SSUSI data.

2.3 Variations with Solar Wind

The final driving sources of auroral variations are the solar wind and the embedded IMF. Especially when the IMF is southward, the magnetic field of the solar wind begins to reconnect to the Earth's magnetic field. Energetic particles carried by the solar wind will

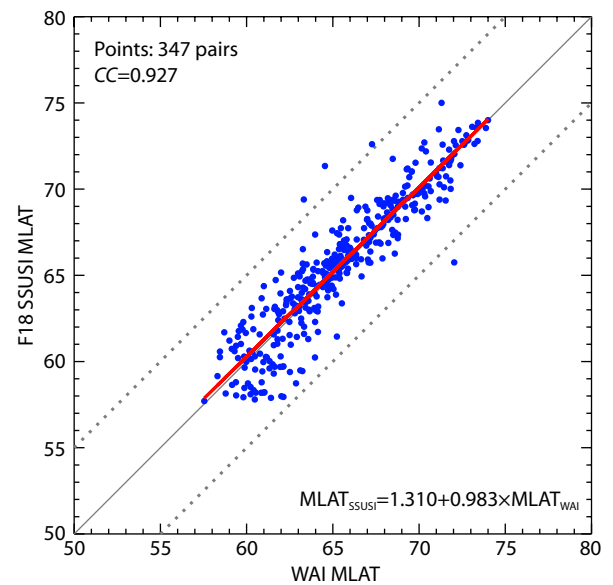


Figure 5. Comparison between auroral equatorward boundaries of WAI and auroral equatorward boundaries of SSUSI. The grey solid line indicates $MLAT_{SSUSI} = MLAT_{WAI}$, and the dashed lines are plotted at $MLAT_{SSUSI} = MLAT_{WAI} \pm 3^\circ$. The red solid lines are the linear fitting of $MLAT_{WAI}$ to $MLAT_{SSUSI}$ with the correlation coefficient shown in the top left corner.

enter the polar upper atmosphere along the Earth's magnetic field lines and then collide with atmospheric particles to trigger the breakup of the aurora.

The variations of solar wind and IMF parameters on 25 August 2018 are shown in Figure 6. It is shown that there was a change of IMF B_z from north to south at about 15:30 (UT). Figure 7 presents the auroral images from WAI before and after the southward turning of the IMF. The solar wind and geomagnetic parameters corresponding to each auroral image are annotated in Figure 6 in different colors.

At 10:18, the B_z was 7.63 nT with a north direction, the solar wind velocity was 404 km/s, and the SYM-H was 15 nT. The auroral image observed by the WAI corresponding to this time is shown in

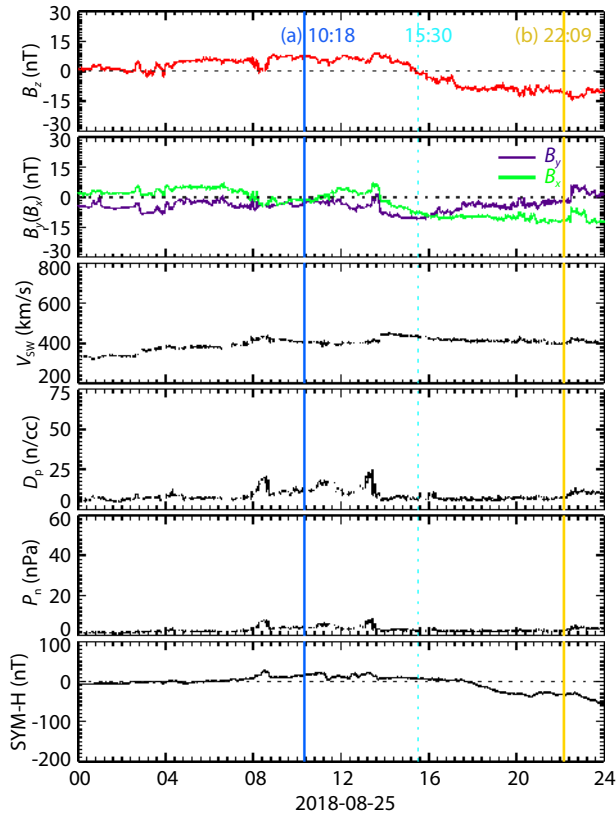


Figure 6. The fluctuation of IMF parameters, solar wind parameters, and SYM-H on 25 August 2018. The IMF B_z turns to be south at 15:30, marked with the dotted line. Before the turning point (a) and after the turning point (b) are chosen to study the evolution of aurora. The auroral images observed by WAI are displayed in Figure 7.

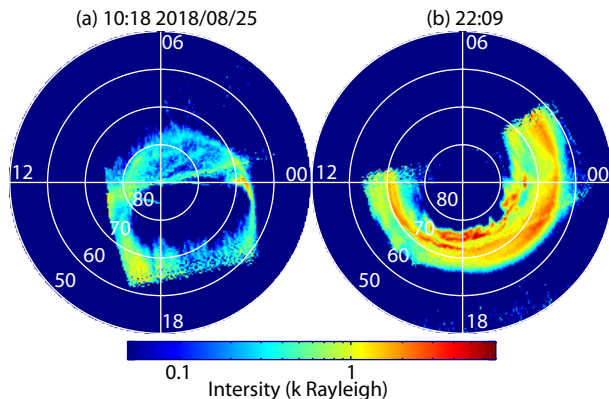


Figure 7. The auroral images observed by the WAI before and after southward turning of IMF. (a) Auroral image obtained at 10:18 UT. (b) Auroral image obtained at 22:09 UT.

Figure 7a. Under the prevailing conditions of quiet solar wind and geomagnetic conditions, the auroral activity was weak, characterized by low emission intensity and an auroral oval of small size.

When the IMF turned to become southward after 15:30 UT, the geomagnetic activity became disturbed, and the SYM-H index began gradually to decrease. At 22:09 UT, the auroral image obtained by the WAI indicates that both the poleward and equator-

ward boundaries of the auroral oval had significantly expanded and the auroral intensity had risen significantly.

3. Results

The WAI onboard the FY-3D can image the aurora with wide FOV, high sensitivity, and high temporal and spatial resolution in FUV wavelengths. The intensity and distribution of the aurora directly reflect geomagnetic activities, the dynamics of the energetic particles near the Polar Regions, and the interaction between the solar wind and the ionosphere/magnetosphere. Based on comparison with the observations of SSUSI/DMSP, it can be concluded that WAI data agree well with SSUSI data in describing both the aurora oval distribution and auroral boundaries. These comparison results indicate that the observations between WAI and SSUSI have fine consistency, and that the WAI data have good quality.

WAI data have thus been used in this paper to investigate the evolution of auroral morphology during geomagnetic events. We have verified that WAI observations can describe in detail the aurora oval evolution with upstream solar wind parameters. We conclude that WAI data can be used effectively to monitor and predict space weather.

4. Conclusions

In this paper, the auroral images of the WAI onboard FY-3D are compared with those of the SSUSI/DMSP. The evolution of auroras with upstream solar wind parameters is also investigated. We report that the WAI data have fine agreement with the SSUSI data, and can depict the aurora morphology during periods of geomagnetic activity. WAI data will play an important role in monitoring space weather and in warning operations at the National Center of Space Weather (National Satellite Meteorological Center) of China, and can also be helpful in study of the coupling of the SW-M-I system.

Acknowledgments

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